

Haverford College

Haverford Scholarship

Faculty Publications

Astronomy

2016

2MTF – V. Cosmography, β , and the residual bulk flow

Karen Masters

Haverford College, klmasters@haverford.edu

Chrisopher M. Springob

Tao Hong

Lister Staveley-Smith

Follow this and additional works at: https://scholarship.haverford.edu/astronomy_facpubs

Repository Citation

Masters, Karen; et al. (2016). "2MTF – V. Cosmography, β , and the residual bulk flow." *Monthly Notices of the Royal Astronomical Society*, 456(2):1886-1900.

This Journal Article is brought to you for free and open access by the Astronomy at Haverford Scholarship. It has been accepted for inclusion in Faculty Publications by an authorized administrator of Haverford Scholarship. For more information, please contact nmedeiro@haverford.edu.

2MTF – V. Cosmography, β , and the residual bulk flow

Christopher M. Springob,^{1,2★} Tao Hong,^{1,2,3} Lister Staveley-Smith,^{1,2}
Karen L. Masters,^{4,5} Lucas M. Macri,⁶ Bärbel S. Koribalski,⁷ D. Heath Jones,⁸
Tom H. Jarrett,⁹ Christina Magoulas⁹ and Pirin Erdoğdu¹⁰

¹International Centre for Radio Astronomy Research, The University of Western Australia, Crawley, WA 6009, Australia

²ARC Centre of Excellence for All-sky Astrophysics (CAASTRO) Crawley, WA 6009, Australia

³National Astronomical Observatories, Chinese Academy of Sciences, 20A Datun Road, Chaoyang District, Beijing 100012, China

⁴Institute for Cosmology and Gravitation, University of Portsmouth, Dennis Sciama Building, Burnaby Road, Portsmouth PO1 3FX, UK

⁵South East Physics Network (www.sepnet.ac.uk) Portsmouth PO1 3FX, UK

⁶George P. and Cynthia Woods Mitchell Institute for Fundamental Physics and Astronomy, Department of Physics and Astronomy, Texas A&M University, 4242 TAMU, College Station, TX 77843, USA

⁷CSIRO Astronomy & Space Science, Australia Telescope National Facility, PO Box 76, Epping, NSW 1710, Australia

⁸Department of Physics and Astronomy, Macquarie University, Sydney, NSW 2109, Australia

⁹Astronomy Department, University of Cape Town, Private Bag X3. Rondebosch 7701, Republic of South Africa

¹⁰Australian College of Kuwait, PO Box 1411, Safat 13015, Kuwait

Accepted 2015 November 9. Received 2015 October 31; in original form 2015 August 17

ABSTRACT

Using the Tully–Fisher relation, we derive peculiar velocities for the 2MASS Tully–Fisher survey and describe the velocity field of the nearby Universe. We use adaptive kernel smoothing to map the velocity field, and compare it to reconstructions based on the redshift space galaxy distributions of the 2MASS Redshift Survey (2MRS) and the *IRAS* Point Source Catalog Redshift Survey (PSCz). With a standard χ^2 minimization fit to the models, we find that the PSCz model provides a better fit to the 2MTF velocity field data than does the 2MRS model, and provides a value of β in greater agreement with literature values. However, when we subtract away the monopole deviation in the velocity zero-point between data and model, the 2MRS model also produces a value of β in agreement with literature values. We also calculate the ‘residual bulk flow’: the component of the bulk flow not accounted for by the models. This is $\sim 250 \text{ km s}^{-1}$ when performing the standard fit, but drops to $\sim 150 \text{ km s}^{-1}$ for both models when the aforementioned monopole offset between data and models is removed. This smaller number is more in line with theoretical expectations, and suggests that the models largely account for the major structures in the nearby Universe responsible for the bulk velocity.

Key words: galaxies: distances and redshifts – galaxies: fundamental parameters – galaxies: spiral – distance scale – large-scale structure of Universe.

1 INTRODUCTION

The velocity field of galaxies exhibits deviations from Hubble flow due to inhomogeneities in the large-scale distribution of matter. By studying the galaxy peculiar velocity field, we can explore the large-scale distribution of matter in the local Universe and so test cosmological models and measure cosmological parameters.

With c representing the speed of light and v_{pec} representing a galaxy’s peculiar velocity, we define the ‘peculiar redshift’ z_{pec} as the peculiar velocity in redshift units, given by

$$z_{\text{pec}} = v_{\text{pec}}/c, \quad (1)$$

where z_{pec} is related to the observed redshift z_{obs} and the redshift due to Hubble flow z_{H} according to

$$(1 + z_{\text{obs}}) = (1 + z_{\text{H}})(1 + z_{\text{pec}}) \quad (2)$$

as given by Harrison (1974). The low-redshift approximation of this relation is

$$v_{\text{pec}} \approx cz_{\text{obs}} - cz_{\text{H}} \approx cz_{\text{obs}} - H_0 D, \quad (3)$$

where H_0 is the Hubble constant and D is the galaxy’s comoving distance.

The measurement of peculiar velocities is thus intertwined with the measurement of distances, in that redshift-independent distances are needed in combination with redshifts to extract peculiar velocities. Most of the largest peculiar velocity surveys have made use

* E-mail: christopher.springob@gmail.com

of one of two redshift-independent distance indicators: the Tully–Fisher relation (TF; Tully & Fisher 1977) and the Fundamental Plane relation (Djorgovski & Davis 1987; Dressler et al. 1987). These scaling relations express the luminosity of a spiral galaxy as a power-law function of its rotational velocity and the radius of an elliptical galaxy as a power-law function of its surface brightness and velocity dispersion, respectively.

The earliest peculiar velocity surveys using these distance indicators, such as Aaronson et al. (1982) and Lynden-Bell et al. (1988), included no more than a few hundred galaxies. Many of these surveys were concatenated into the Mark III catalogue (Willick et al. 1995; Willick et al. 1996). Among the earliest individual TF surveys to include more than ~ 1000 galaxies were a set of surveys conducted by Giovanelli, Haynes, and collaborators (e.g. Giovanelli et al. 1994, 1995, 1997a; Haynes et al. 1999a,b). These surveys were combined, along with additional data, to create the SFI++ survey (Masters et al. 2006; Springob et al. 2007), which included TF data for ~ 5000 galaxies.

At the time of its release, SFI++ was the largest peculiar velocity survey compiled. However, both its selection criteria and data sources were quite heterogeneous. The 2MASS Tully–Fisher survey (2MTF) was envisioned as a ‘cleaner’ all-sky TF survey, with more stringent selection criteria, drawn from a more homogeneous data set, and extending to significantly lower Galactic latitudes. It draws on galaxies selected from the 2MASS Redshift Survey (2MRS; Huchra et al. 2012), and uses photometry from the 2MASS extended source catalogue (Jarrett et al. 2000) and spectroscopy from the Green Bank Telescope (GBT), Parkes radio telescope, Arecibo telescope, and other archival H I catalogues. The archival data overlaps heavily with the archival data set used for SFI++, and the H I width measurement procedure and peculiar velocity derivation also follows the SFI++ procedure closely. The 2MTF template relation was presented by Masters, Springob & Huchra (2008), while the first cosmological analysis using the data set, a measurement of the bulk flow, was presented by Hong et al. (2014).

In this paper, we examine the cosmography of the observed 2MTF velocity field, and compare it to two reconstructions of the predicted peculiar velocity field, which assume that the matter distribution traces the galaxy distribution. The first is the Erdoğan et al. (submitted, updated from Erdoğan et al. 2006) reconstruction of the aforementioned 2MRS. The second is the Branchini et al. (2001) reconstruction of the IRAS Point Source Catalog Redshift Survey (PSCz; Saunders et al. 2000). In comparing to these reconstructions, we compute the χ^2 agreement between the observed and predicted velocity fields, the value of the redshift space distortion parameter β , and the amplitude and direction of the ‘residual bulk flow’, which represents the component of the bulk flow of the local Universe not predicted by models.

The history of such comparisons between observed peculiar velocity fields and predictions made from large all-sky redshift surveys goes back to analyses such as those of Kaiser et al. (1991), Shaya, Tully & Pierce (1992), Hudson (1994), Davis, Nusser & Willick (1996), and Hudson et al. (2004). More recently, we have seen the comparisons between SFI++ and 2MRS performed by Davis et al. (2011), between SFI++ and 2M++ (Lavaux & Hudson 2011) performed by Carrick et al. (2015), and between the 6dF Galaxy Survey (6dFGS) peculiar velocity field (Springob et al. 2014) and both the 2MRS and PSCz reconstructions performed by Magoulas (2012), Magoulas et al. (in preparation), and Springob et al. (2014). These later papers find good overall agreement between the observed and predicted velocity fields, but disagree on whether the predicted fields can explain the amplitude of the bulk flow. 2MTF

is well suited to examine this issue because of its all-sky coverage, sampling the sky down to Galactic latitudes of $|b| = 5^\circ$.

The paper is arranged as follows: In Section 2, we describe the 2MTF observational data set as well as the 2MRS and PSCz model velocity fields. In Section 3, we describe the derivation of peculiar velocities. In Section 4, we describe the cosmography of the 2MTF velocity field, present our comparison of the field to the model velocity fields, and discuss our results. We summarize our results in Section 5.

2 DATA

2.1 2MTF TF data

The 2MTF target list was compiled from the set of all 2MRS (Huchra et al. 2012) spirals with total K -band magnitude $K_s < 11.25$, cosmic microwave background (CMB) frame redshift $cz < 10\,000$ km s $^{-1}$, and axis ratio $b/a < 0.5$. There are ~ 6000 galaxies that meet these criteria, though many of them are quite faint in H I, and observationally expensive to observe with the available single dish radio telescopes. Thus only a fraction of these objects are included in the final 2MTF sample.

We have combined archival data with observations made as part of the Arecibo Fast Legacy ALFA Survey (ALFALFA; Giovanelli et al. 2005), and observations made with the GBT (Masters et al. 2014b) and Parkes telescope (Hong et al. 2013). While the new GBT and Parkes observations preferentially targeted late-type spirals, the archival data includes all spiral types. The complete 2MTF sample includes 2018 galaxies (note that this is separate from the 888 cluster galaxies used to fit the template TF relation in Masters et al. 2008). However, as noted in Section 3.2, for the analysis performed in this paper, we exclude large outliers from the TF relation, which reduces the number of objects to 1985.

2.1.1 Photometric data

All of our photometry is drawn from the 2MRS catalogue. 2MRS is an all-sky redshift survey, consisting of $\sim 43\,000$ redshifts of 2MASS galaxies, extending in magnitude to $K_s < 11.75$ and Galactic latitude $|b| > 5^\circ$. While the final sample (Huchra et al. 2012) has a limiting K -band magnitude of 11.75, the 2MTF survey began before 2MRS was complete, and so the magnitude limit was set to the somewhat shallower value of 11.25. For the analysis presented in this paper, we use total K -band magnitudes.

While the Masters et al. (2008) template used I -band and J -band axis ratios, we use 2MASS $J/H/K$ co-added axis ratios (see Jarrett et al. 2000). As explained by Hong et al. (2014), the dispersion between these two definitions of axis ratio is ~ 0.096 , and we account for this in deriving the scatter of the TF relation, but find no systematic trend between the two definitions towards larger or smaller values.

Internal dust extinction and k -correction were done as described by Masters et al. (2008), and updated in the erratum Masters, Springob & Huchra (2014a). The 2MRS total magnitudes are already corrected for Galactic extinction, so no further correction is necessary.

2.1.2 Spectroscopic data

The primary source of archival data used here is the Cornell H I archive (Springob et al. 2005), which offers H I spectroscopy for

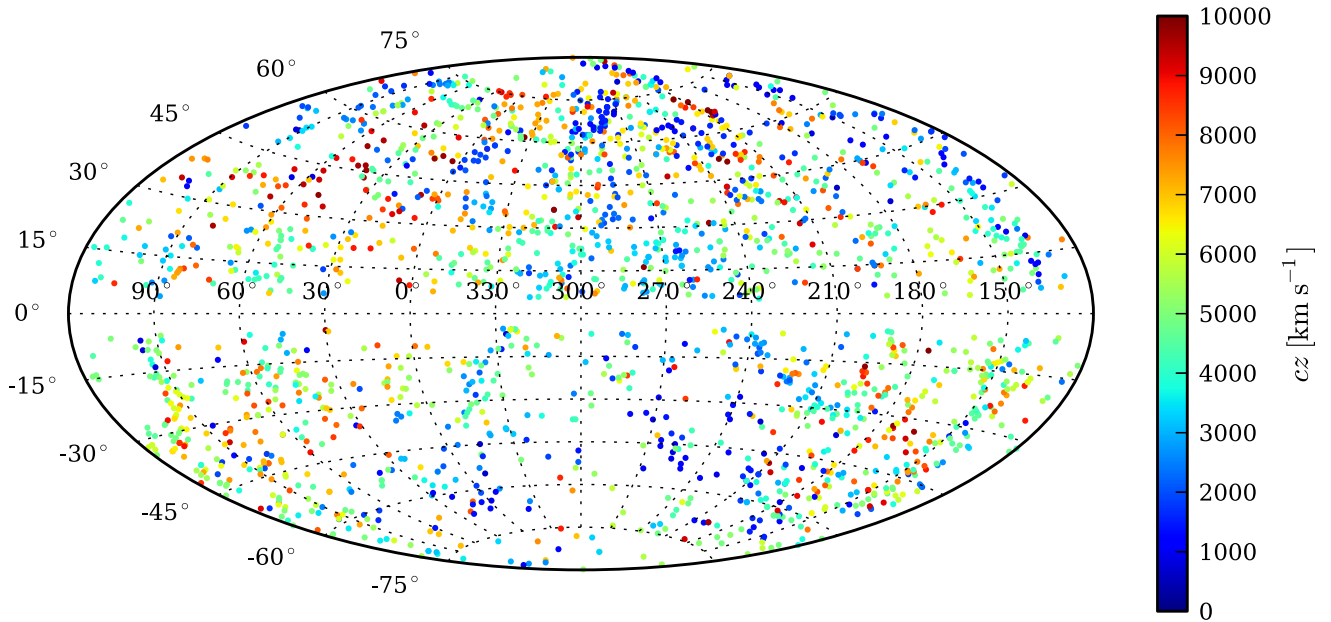


Figure 1. Distribution of 2MTF galaxies in Galactic latitude (l) and longitude (b), shown in an Aitoff projection. Galaxies are colour-coded by redshift, according to the colour bar on the right-hand side of the plot.

~ 9000 galaxies in the local Universe, as observed by single dish radio telescopes. We find 1038 galaxies from this data set with high-quality spectra that match the 2MTF selection criteria. We also include H I data from Theureau et al. (1998), Theureau et al. (2005), Theureau et al. (2007), and Matthewson, Ford & Buchhorn (1992), as well as the Nancay observations from table A.1 of Paturel et al. (2003). The raw observed H I widths drawn from these sources were then corrected for inclination, redshift stretch, instrumental effects, turbulence, and smoothing, according to the prescriptions of Springob et al. (2005), which were then updated by Hong et al. (2013).

To supplement the archival data, we made new observations using the Parkes and GBT between 2006 and 2012. These data sets were presented by Hong et al. (2013) and Masters et al. (2014b), respectively, but we briefly summarize them here.

The observations were divided up in such a way that Parkes was used to target galaxies in southern declinations ($\delta < -40^\circ$), while the GBT targeted more northern galaxies ($\delta > -40^\circ$), but outside the ALFALFA survey region. (See Fig. 1 for the sky distribution of objects.) The GBT observations targeted 1193 galaxies in position switched mode, with the spectrometer set at 9 level sampling with 8192 channels. After smoothing, the velocity resolution was 5.15 km s^{-1} . 727 galaxies were detected, with 483 of them being deemed sufficiently high-quality detections to be included in our sample.

For the Parkes observations, we targeted 305 galaxies which did not already have high-quality H I width measurements in the literature. Of these, we obtained width measurements suitable for inclusion in our sample for 152 galaxies. The multibeam correlator produced raw spectra with a velocity resolution of 1.6 km s^{-1} , though Hanning smoothing broadened the resolution to 3.3 km s^{-1} . Each of the GBT and Parkes spectra were analysed using the IDL routine *awv_fit.pro*, which is based on the method used for the Cornell H I archive (Springob et al. 2005), which in turn is based on the earlier approach developed by Giovanelli, Haynes, and collaborators (e.g. Giovanelli et al. 1997b). We use the W_{F50} width algorithm, as defined in those papers. It involves fitting a line to either side of

the H I line profile, and measuring the width from the points on each line representing 50 per cent of the flux minus *rms* value. This approach was also adopted for use in the ALFALFA survey.

In addition to the archival, GBT, and Parkes data sets, we also included ALFALFA data from the initial data release (Haynes et al. 2011). The catalogue presented in that data release covers roughly 40 per cent of the final survey. However, the ALFALFA team has provided us with additional unpublished data from the survey, current as of 2013 October. In total then, we cover ~ 66 per cent of the ALFALFA sky. From this sample, we use 576 galaxy widths, which will be updated once the final release of ALFALFA is available.

As noted above, this gives us a final sample of 1985 galaxies when TF outliers are excluded. The redshift histogram of this sample can be found in Fig. 2.

2.2 PSCz model velocity field

The *IRAS* PSCz (Saunders et al. 2000) includes 15 500 galaxies, covering 84 per cent of the sky, with most of the missing sky area lying at low Galactic latitudes. Branchini et al. (1999) reconstructs the local density and velocity field from this survey, using a spherical harmonic expansion method proposed by Nusser & Davis (1994). The method assumes a linear mapping between the PSCz spatial distribution of galaxies and the matter distribution, with an input assumed bias parameter $\beta = 0.5$. The grid spacing of the model is $2.8 h^{-1} \text{ Mpc}$, extending to a distance of $180 h^{-1} \text{ Mpc}$ from the origin in each direction.

We convert the PSCz velocity grid from real space to redshift space, so that we may compare to the redshift space positions of the 2MTF galaxies. Each gridpoint is assigned to its position in redshift space by adding its reconstructed peculiar velocity to its real distance. The reconstructed velocities are then linearly interpolated on to a regularly spaced grid (again with $2.8 h^{-1} \text{ Mpc}$ resolution) in redshift space. The problem of triple-valued regions is mitigated in the original Branchini et al. (1999) reconstruction, by the authors collapsing galaxies within clusters, and applying a method devised

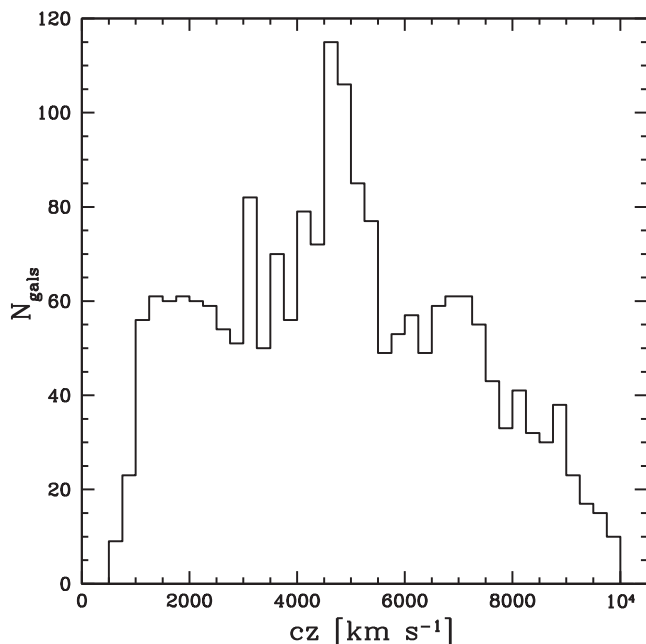


Figure 2. Redshift distribution of galaxies in 2MTF in the CMB reference frame. The bin width is 250 km s^{-1} .

by Yahil et al. (1991) to determine the locations of galaxies along those lines of sight.

Following the reconstruction of the PSCz velocities on to this redshift space grid, we assign line-of-sight predicted PSCz velocities to each of the galaxies in the 2MTF sample. This is done by interpolating the nearest gridpoints to the position of the galaxy in question. The interpolation is done with inverse distance weighting. This method was also used by Springob et al. (2014) to match 6dFGS peculiar velocities to predictions from models.

2.3 2MRS model velocity field

We also make use of a density/velocity field reconstruction from 2MRS itself. As noted in Section 2.1, the final 2MRS data release (Huchra et al. 2012) includes redshifts for 44 699 galaxies, and covers 91 per cent of the sky (excluding only the Galactic zone of avoidance). Erdođdu et al. (2006) reconstructed the density/velocity field from an earlier 2MRS data release (Huchra et al. 2005), and this has now been updated to make use of the final data set (Erdođdu et al. submitted).

The reconstruction method is described in detail by Erdođdu et al. (2006). It most closely follows the method used by Fisher et al. (1995), again assuming that the matter distribution follows the galaxy distribution, with an assumed $\beta = 0.4$. It involves decomposing the redshift space density field of galaxies into spherical harmonics, and smoothing with a Wiener filter. The reconstruction gives densities and velocities on a grid in supergalactic Cartesian coordinates with gridpoints spaced by $8 h^{-1}$ Mpc and extending to a distance of $200 h^{-1}$ Mpc from the origin in each direction.

As with PSCz, we convert this grid from real space to redshift space. The resulting redshift space grid in this case has grid spacing of $4 h^{-1}$ Mpc. Again, as with PSCz, we then use interpolation between gridpoints to assign model velocities to galaxies in the 2MTF sample.

3 TF DISTANCES AND PECULIAR VELOCITIES

We follow the same basic procedure that was followed for the SFI++ derivation of peculiar velocities by Springob et al. (2007), which relied on the calibration of the TF relation from the template relation of Masters et al. (2006). This in turn followed the procedure used by the SFI and SCI surveys (Giovanelli et al. 1994, 1995, 1997b).

In brief, we use the J -, H -, and K -band TF template relations from Masters et al. (2008) to derive the peculiar velocities of the individual galaxies in our sample. The Masters et al. (2008) template sample includes 888 galaxies, while the sample presented in this paper includes 1985 galaxies which does not overlap with the template sample. The template relations found by Masters et al. (2008) are

$$M_K - 5 \log h = -22.188 - 10.74(\log W - 2.5),$$

$$M_H - 5 \log h = -21.951 - 10.65(\log W - 2.5),$$

$$M_J - 5 \log h = -21.370 - 10.61(\log W - 2.5), \quad (4)$$

where W is the corrected H I width in units of km s^{-1} , and M_K , M_H , and M_J are the corrected absolute magnitudes in the three bands.

One difference between the SFI++ approach and the one employed here is that we work with logarithmic distance ratios throughout, rather than converting to linear peculiar velocities. This is done because the distance errors (as well as the individual errors on line width and magnitude) are approximately lognormal. We thus make use of the quantity

$$\Delta d^* = \log \left(\frac{d_z}{d_{\text{TF}}^*} \right) = \frac{-\Delta M}{5}, \quad (5)$$

where $\Delta M = M_{\text{obs}} - M(W)$ is the difference between the corrected absolute magnitude M_{obs} , calculated using the redshift distance of the galaxy (d_z), and the magnitude $M(W)$ derived from the TF template relation. d_{TF}^* is the distance to the galaxy derived from the TF relation, but not corrected for Malmquist/selection bias. We then refer to Δd^* as the logarithmic distance ratio (uncorrected for Malmquist bias). In Section 3.2, we discuss the Malmquist bias correction, at which point this quantity is replaced by Δd , the logarithmic distance ratio (corrected for Malmquist bias).

As noted by Hong et al. (2014), the fact that Masters et al. (2008) uses a different set of axial ratios means that we must derive new measurements of the intrinsic scatter in the TF relation. Hong et al. (2014) does this, and arrives at the relations

$$\epsilon_{\text{int},K} = 0.44 - 0.66(\log W - 2.5),$$

$$\epsilon_{\text{int},H} = 0.44 - 0.95(\log W - 2.5),$$

$$\epsilon_{\text{int},J} = 0.46 - 0.77(\log W - 2.5), \quad (6)$$

for the scatter in the TF relation *in magnitude units*. To convert these to logarithmic distance units, one must divide the ϵ_{int} values by 5.

3.1 Galaxy groups

Crook et al. (2007) describes a galaxy group catalogue for 2MRS, derived using a ‘friends-of-friends’ algorithm. 55 of these groups include more than one member in our sample. To eliminate the effects of motions within galaxy groups, we thus fix the redshift used when calculating the logarithmic distance ratio in equation (5) to the group redshift for all of our group galaxies. However, the magnitude offset ΔM is still calculated separately for each galaxy.

3.2 Selection bias

‘Malmquist bias’ is the term used to describe biases originating from the interaction between the spatial distribution of objects and the selection effects (Malmquist 1924). It results from the coupling between the random distance errors and the apparent density distribution along the line of sight. There are two types of distance errors to consider. The first is ‘inhomogeneous Malmquist bias’, which arises from local density variations due to large-scale structure along the line of sight. This bias is most pronounced when measuring galaxy distances in real space. This is because the large distance errors scatter the measured galaxy distances away from overdense regions, creating artificially inflated measurements of infall on to large structures. By contrast, when the measurement is done in redshift space, the much smaller redshift errors mean that this effect tends to be negligible (see e.g. Strauss & Willick 1995).

For the 2MTF sample, we are measuring galaxy distances and peculiar velocities in redshift space rather than real space. In this case, inhomogeneous Malmquist bias is negligible, and the form of Malmquist bias that we must deal with is of the second type, known as homogeneous Malmquist bias, which affects all galaxies independently of their position on the sky. It is a consequence of both (1) the volume effect, which means that more volume is covered within a given solid angle at larger distances than at smaller distances, and (2) the selection effects, which cause galaxies of different luminosities, radii, velocity dispersions etc. to be observed with diminishing completeness with increased distance. We note, however, that different authors use somewhat different terminology, and the latter effect described above is often simply described as ‘selection bias’.

The approach one takes in correcting for this bias depends in part on the selection effects of the survey. In our case, we used homogeneous criteria in determining which galaxies to observe. However, many of the galaxies that met our selection criteria yielded non-detections, marginal detections, or there was some other problem with the spectrum that precluded the galaxy’s inclusion in our sample. Our final sample, then, lacks the same homogeneity as the original target sample, and our selection bias correction procedure must account for this.

We adopt the following procedure.

(1) Using the stepwise maximum likelihood method (Efstathiou, Ellis & Peterson 1988), we derive the K -band luminosity function $\Phi(M_k)$, as a function of K -band absolute magnitude M_k , for all galaxies in 2MRS that meet our K -band apparent magnitude, Galactic latitude, morphological, and axis ratio criteria. For this purpose, we include galaxies beyond the 10 000 km s⁻¹ redshift limit, to simplify the implementation of the luminosity function derivation. The resulting luminosity function of this sample is found in Fig. 3. We fit a Schechter function (Press & Schechter 1974) to this distribution, and find parameters $M_k^* = -23.1$ and $\alpha = -1.10$. The final parameter of the Schechter function is the normalization. The stepwise maximum likelihood method cannot derive this parameter, but it is not needed for our method in any case. This luminosity function has a steeper slope than the 2MASS K -band luminosity function derived by Kochanek et al. (2001), who find $\alpha = -0.87$. Our morphological selection criteria presumably contribute to this difference, though Jones et al. (2006) notes that there is some disagreement between the Kochanek et al. (2001) luminosity function and others.

(2) We define the ‘completeness’ as the fraction of the target sample that is included in our final catalogue for a given apparent magnitude bin. We divide up the sky between two regions: one

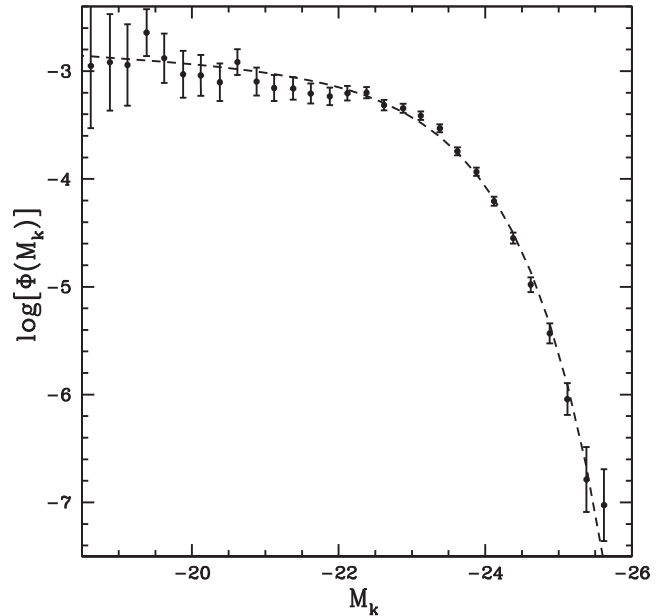


Figure 3. K -band luminosity function $\Phi(M_k)$, as a function of absolute magnitude M_k , for 2MTF target galaxies, derived using the stepwise maximum likelihood method. Error bars represent the uncertainty from Poisson statistics. The dashed line represents the best-fitting Schechter function, which gives Schechter parameters $M_k^* = -23.1$ and $\alpha = -1.10$. The vertical axis is in dex units, but the normalization is arbitrary, as the stepwise maximum likelihood method is unable to fit the normalization. However, the normalization is also irrelevant, as the correction for selection bias only requires us to know the relative luminosity function.

covering all declinations north of $\delta = -40^\circ$, and the other covering all declinations south of $\delta = -40^\circ$. The $\delta = -40^\circ$ boundary is the declination at which the sample transitions from GBT to Parkes observations. After experimenting with further subdivisions of the sky, we conclude that the completeness as a function of apparent magnitude remains constant within each region, and assume that that holds in the analysis that follows. We compute this completeness function separately for the two sky regions, by taking the ratio of observed galaxies to galaxies in the target sample for K -band apparent magnitude bins of width 0.25 mag.

(3) For each of the galaxies in the sample, we compute the uncorrected logarithmic distance ratio (Δd^*) probability distribution, assuming a Gaussian distribution, with an uncorrected 1σ scatter of ϵ_{d^*} . For each logarithmic distance ratio value (Δd_i) within 2σ of the nominal value of Δd^* , we weight the probability by w_i , where $1/w_i$ is the completeness (from Step 2) integrated across the K -band luminosity function (from Step 1), evaluated at the Δd_i in question. Using the luminosity function in this manner involves converting it into apparent magnitudes, using the relevant distance modulus.

(4) Finally, we fit a Gaussian to the re-weighted probability distribution. This gives us the corrected logarithmic distance ratio (Δd), and its error (ϵ_d). We then cut from the sample any galaxies deemed outliers, defined as those for which the deviation from $\Delta d = 0$ cannot be accounted for by the quadrature sum of a 300 km s⁻¹ scatter in peculiar velocities plus 3σ deviation from the TF relation. 43 galaxies are eliminated by this criterion, leaving us with a total sample of 1985 galaxies.

The Δd histograms can be found in Hong et al. (2014). As noted in that paper, the mean TF distance error translates to ~ 22 per cent in all three wavebands.

4 RESULTS AND DISCUSSION

4.1 Velocity field cosmography

Because of the significant distance errors obtained for each individual galaxy, we use adaptive kernel smoothing to get a cosmographic view of the velocity field. We set up a 3D redshift space grid in supergalactic Cartesian coordinates, with all gridpoints spaced $4 h^{-1}$ Mpc apart. At each gridpoint, we compute adaptively smoothed values of Δd for both the PSCz reconstructed velocity field, and the observed 2MTF field. This is done following a procedure outlined in Springob et al. (2014), that draws on methods used by Silverman (1986) and Ebeling, White & Rangarajan (2006). We summarize this method below.

We define $\Delta d(\mathbf{r}_i)$ as the logarithmic distance ratio at redshift space position \mathbf{r}_i . We aim to recover $\Delta d(\mathbf{r}_i)$ by smoothing the individual logarithmic distance ratios Δd_j for galaxy j . Our smoothing algorithm is defined as

$$\Delta d(\mathbf{r}_i) = \frac{\sum_{j=1}^{N_j} \Delta d_j \cos \theta_{i,j} e^{-rr_{i,j}/2} \sigma_j^{-3}}{\sum_{j=1}^{N_j} e^{-rr_{i,j}/2} \sigma_j^{-3}}, \quad (7)$$

where σ_j is the smoothing length of the 3D Gaussian kernel for galaxy j , $\theta_{i,j}$ is the angle between the \mathbf{r}_i for gridpoint i and the \mathbf{r}_j for galaxy j , and $rr_{i,j}$ is the square of the distance between gridpoint i and galaxy j in units of σ_j . The summations in both the numerator and denominator of equation (7) run over all N_j galaxies for which $rr_{i,j} < 9$.

We define the smoothing length σ_j as a function of the fiducial kernel $\sigma' = 15 h^{-1}$ Mpc, weighted as a function of local density, δ_j :

$$\sigma_j = \sigma' \left[\frac{\exp(\sum_{l=1}^N \ln \delta_l / N)}{\delta_j} \right]^{1/2}, \quad (8)$$

where

$$\delta_j = \sum_{k=1}^{N_k} e^{-rr_{j,k}/2} \quad (9)$$

and $rr_{j,k}$ is the square of the distance between galaxy j and galaxy k in units of σ' . The sum in equation (8) is over all N galaxies in the survey, while the sum in equation (9) is over the N_k galaxies within $3\sigma'$ of galaxy j . In this case, we set $\sigma' = 15 h^{-1}$ Mpc. The mean smoothing length is $\langle \sigma_j \rangle = 7.2 h^{-1}$ Mpc, with a 1σ scatter of $4.2 h^{-1}$ Mpc. The smallest smoothing length is $2.4 h^{-1}$ Mpc, while the largest is $29.0 h^{-1}$ Mpc.

The distribution of smoothing lengths only weakly depends on the fiducial value, σ' . The selection of the fiducial value is somewhat arbitrary, but we chose a value that we find clearly illustrate the broad trends in the velocity field. In any case, the smoothed map is not used for any quantitative analysis in this paper. All measurements of β and the residual bulk flow use the unsmoothed Δd values from the individual galaxies.

In Figs 4–6, we show the adaptively smoothed velocity field along slices of SGX, SGY, and SGZ, for both the observed 2MTF data set and the 2MRS and PSCz models. SGX, SGY, and SGZ are the orthogonal axes in supergalactic Cartesian coordinates. They are defined by $SGX = r \cos(sgb) \cos(sgl)$, $SGY = r \cos(sgb) \sin(sgl)$, $SGZ = r \sin(sgb)$, where r is the redshift space distance to the galaxy, and sgl and sgb are the supergalactic longitude and latitude, respectively.

In these figures, we also provide plots of the residual logarithmic distance ratios, $\Delta d_{2MTF} - \Delta d_{2MRS}$ and $\Delta d_{2MTF} - \Delta d_{PSCz}$. These

plots indicate features of the velocity field which are not predicted by the models. We also include the approximate positions of various nearby features of large-scale structure. (It should be noted that though objects such as the Hydra-Centaurus Supercluster and the Pisces-Perseus Supercluster are extended structures, we only display them as point sources so as not to obscure other features of the maps.)

The most notable features distinguishing the observed velocity field from the model velocity fields are: (1) a monopole deviation in Δd between the 2MRS reconstruction and both the PSCz reconstruction and the 2MTF observation, with the former showing many more gridpoints with positive values of Δd , and (2) a dipole deviation between the observed velocity field and both model velocity fields. Namely, the observed velocity field shows a noticeably large motion towards negative SGX and positive SGY, more or less in the direction of the Hydra-Centaurus and Shapley Superclusters. Hong et al. (2014) measures this bulk flow at depths of 20, 30, and $40 h^{-1}$ Mpc with a Gaussian window function, finding values of 310.9 ± 33.9 , 280.8 ± 25.0 , and $292.3 \pm 27.8 \text{ km s}^{-1}$, respectively. While Hong et al. (2014) found these bulk flow values to be in good agreement with expectations from the Λ cold dark matter (Λ CDM) model, we see here that the observed bulk flow does not appear to be replicated in the velocity field reconstructions. We discuss this issue greater detail in Section 4.3.

Figs 4–6 can also be directly compared to previous maps of the velocity field. For example, Theureau et al. (2007) presents similar cosmographic plots of the nearby velocity field based on the Kinematics of the Local Universe survey (KLUN; Theureau et al. 2005), as do Tully et al. (2014) and Hoffman, Courteau & Tully (2015) using the Cosmic Flows 2 survey (Tully et al. 2013). Springob et al. (2014) also includes very similar figures, as derived from the 6dFGS velocity field. Theureau et al. (2007) fig. 11 closely matches the ‘mid-plane’ panel in Fig. 4 from this paper, for example. In both cases, we see evidence of both foreground and background infall on to both the Pisces-Perseus Supercluster and Coma. Both also show foreground infall towards Hydra-Centaurus, though our 2MTF cosmography plots are not deep enough to establish whether there is backside infall as seen by Theureau et al. (2007). Again, though, the largest apparent difference between the 2MTF and KLUN velocity field maps would appear to be the large dipole apparent in 2MTF, which is also notable in the 6dFGS cosmography, as described by Springob et al. (2014).

4.2 Fitting to the model velocity fields and measuring β

As noted in Sections 2.2 and 2.3, the fiducial values of the redshift space distortion parameter β are 0.4 for the 2MRS model and 0.5 for the PSCz model. These are values that are assumed by the models, but we can measure the values for each model directly using 2MTF. To clarify, β is related to the matter density of the Universe (in units of the critical density) Ω_m and the linear bias parameter b , according to

$$\beta = \Omega_m^{0.55} / b \quad (10)$$

in a flat Λ CDM universe (Linder 2005). b is defined as $b = \delta_g / \delta_m$, the ratio between overdensities in the galaxy density field and overdensities in the matter density field. In the linear regime, the induced velocity v at position r is then

$$v(r) = \frac{\beta}{4\pi} \int d^3 r' \frac{r' - r}{|r' - r|^3} \delta_g(r'). \quad (11)$$

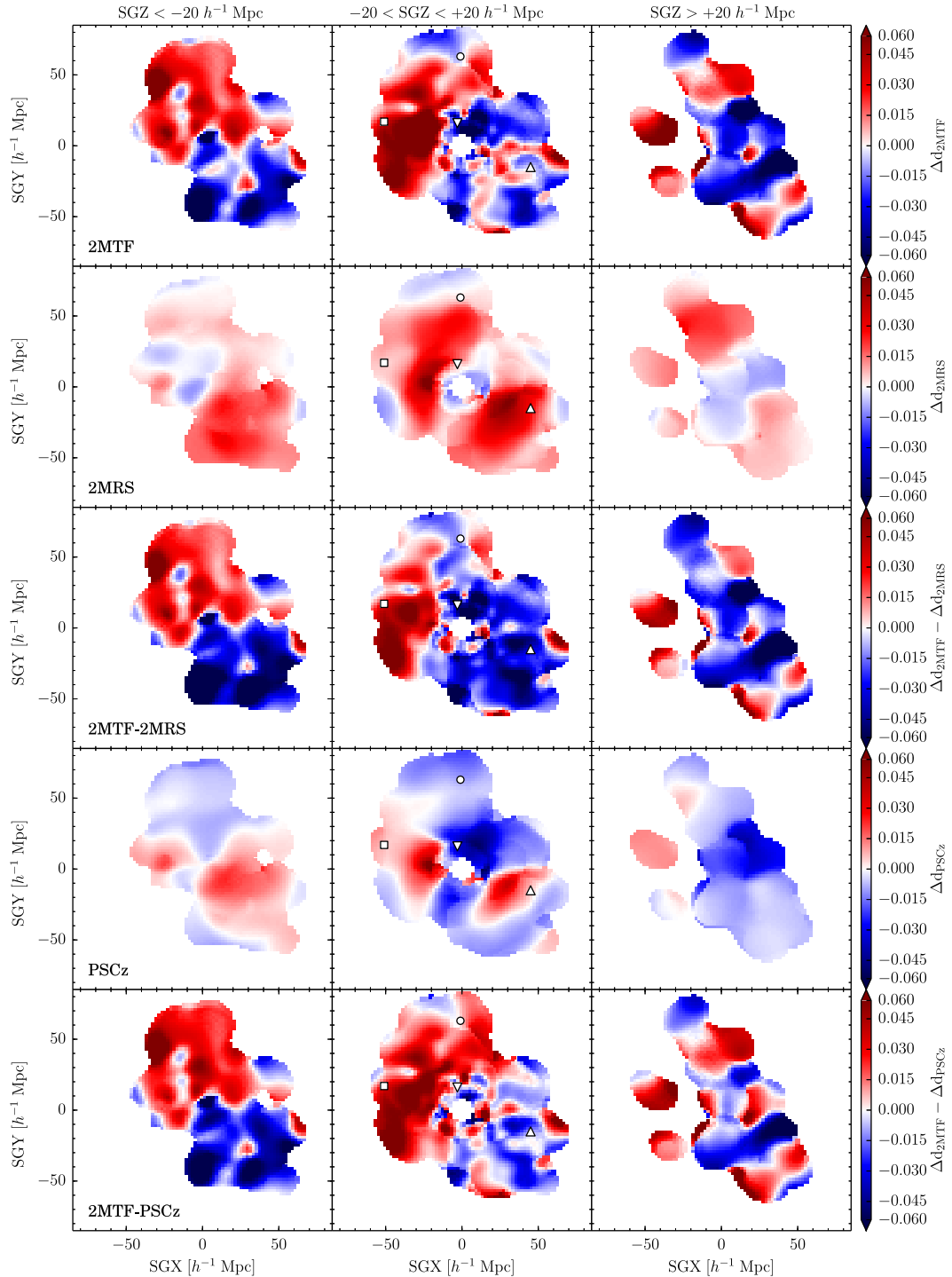


Figure 4. Adaptively smoothed maps of the nearby galaxy velocity field in supergalactic Cartesian coordinates, in slices of SGZ. In each case, the velocity field is given in logarithmic distance units ($\Delta d = \log(D_z/D_H)$), in the nomenclature of Section 3, as the logarithm of the ratio between the redshift distance and the true Hubble distance. As shown in the colour bars for each panel, redder (bluer) colours correspond to more positive (negative) values of the logarithmic distance ratio, Δd , and thus more positive (negative) peculiar velocities. The left-hand column is the slice for $SGZ < -20 h^{-1} \text{ Mpc}$. The middle column is the $-20 < SGZ < +20 h^{-1} \text{ Mpc}$ slice. The right-hand column is the $SGZ > +20 h^{-1} \text{ Mpc}$ slice. The top row is the observed velocity field from 2MTF. The second row is the 2MRS reconstructed velocity field, as derived by Erdoğan et al. (submitted). The third row is the difference between the observed 2MTF Δd values and the predicted 2MRS Δd values. That is, it is the residual in Δd from 2MRS. The fourth row is the PSCz reconstructed velocity field, as derived by Branchini et al. (1999). The fifth row is the 2MTF residual from PSCz. Gridpoints are spaced $4 h^{-1} \text{ Mpc}$ apart. We also mark the approximate positions of several features of large-scale structure: The Coma Cluster (○), the Hydra-Centaurus Supercluster (□), the Virgo Cluster (▽), and the Pisces-Perseus Supercluster (Δ).

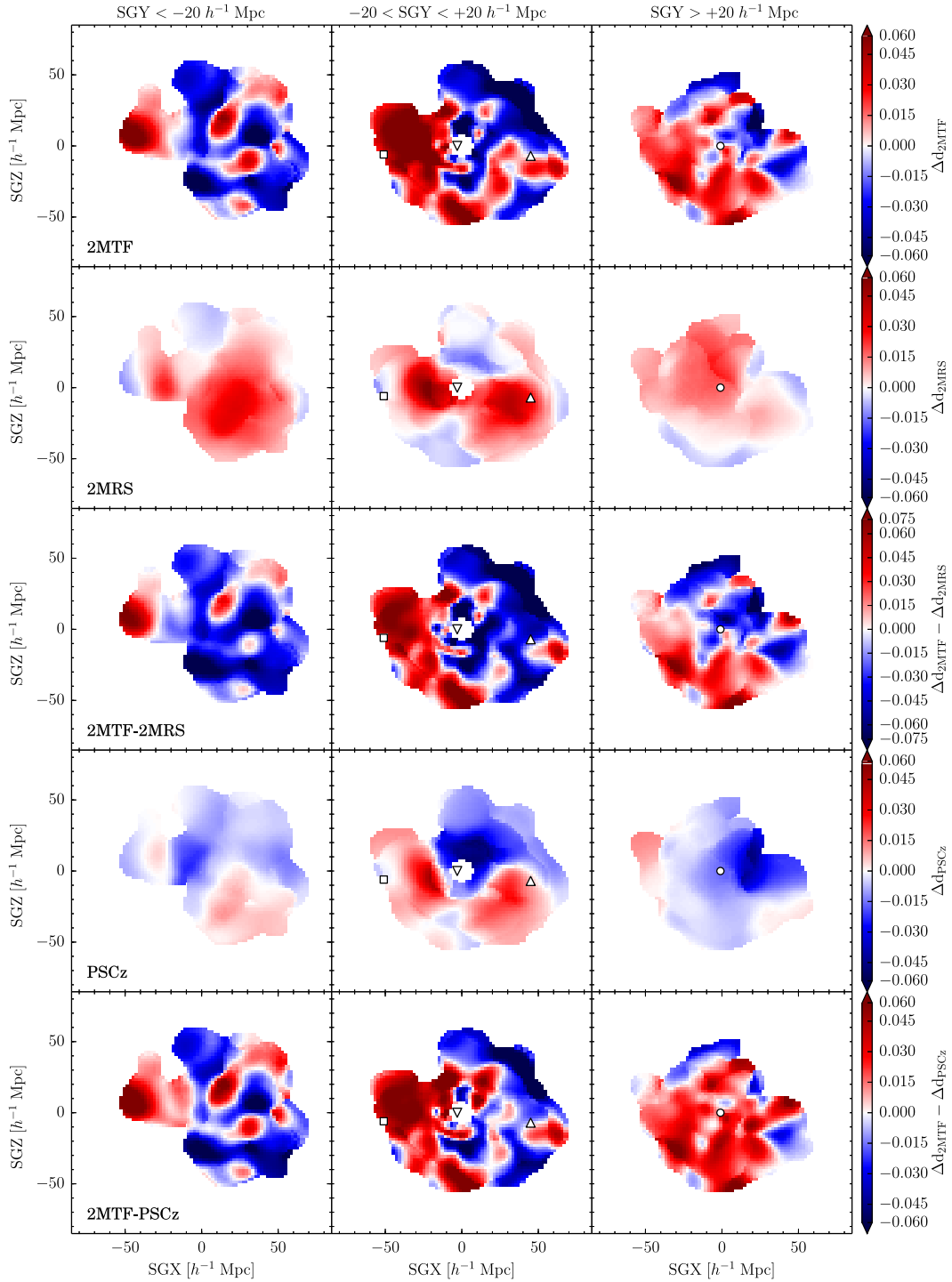


Figure 5. Same as Fig. 4, but showing slices parallel to the SGY axis rather than the SGZ axis.

Since the amplitude of velocities is proportional to the value of β , the characteristic observed amplitude should scale with β . Thus

$$|v_{\text{obs}}|/|v_{\text{model}}| = \beta/\beta_{\text{fid}}. \quad (12)$$

Here, we employ a simple χ^2 minimization in order to measure the value of β for both the 2MRS and PSCz models, using the N

galaxies in the 2MTF sample:

$$\chi_v^2 = \sum_{i=1}^N \frac{(\Delta d_i - \Delta d_{\text{model},i})^2}{N\sigma_i^2}, \quad (13)$$

where Δd_i is the logarithmic distance ratio of galaxy i , and σ_i is the uncertainty on this quantity. $\Delta d_{\text{model},i}$ is the logarithmic distance

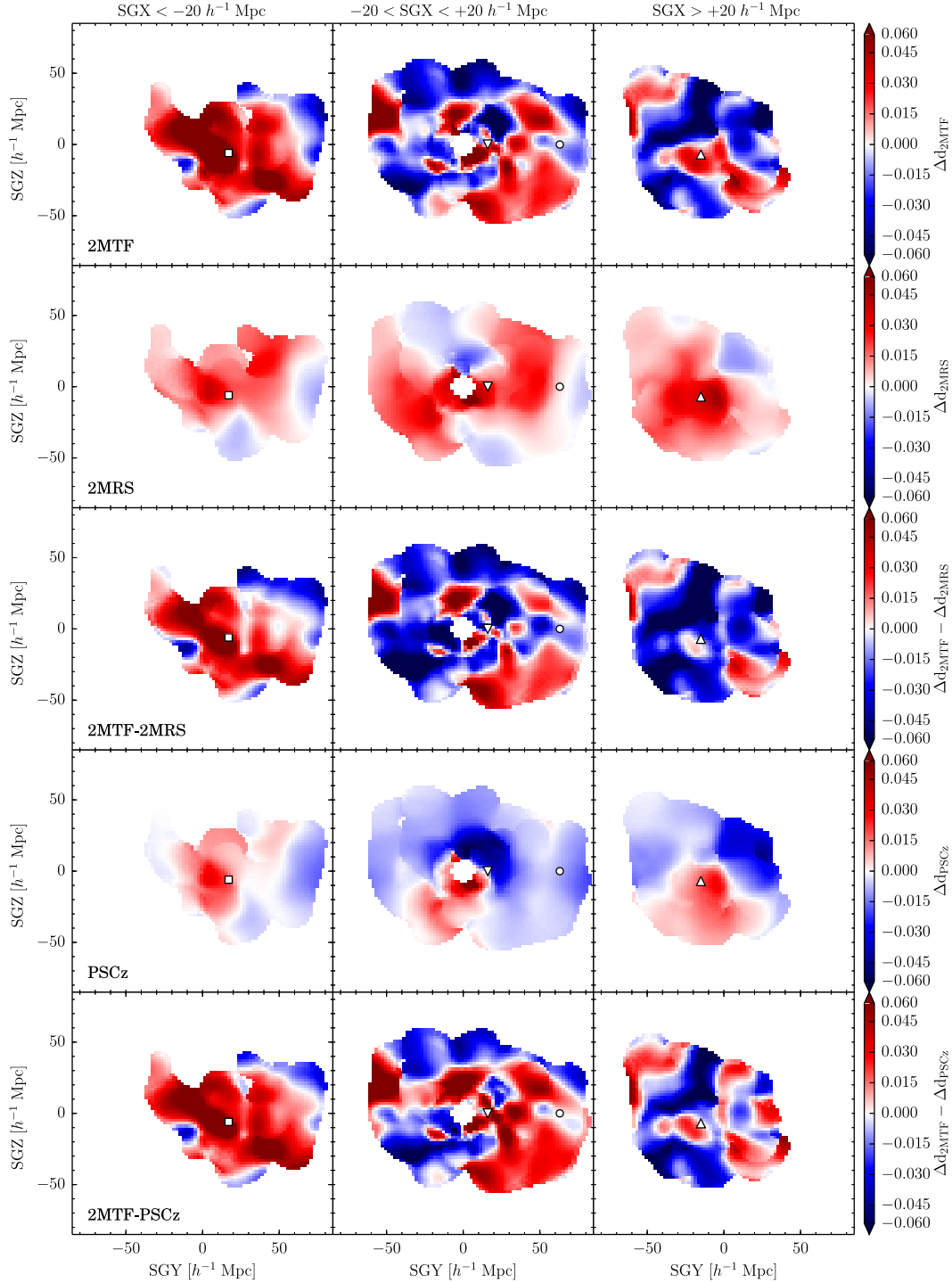


Figure 6. Same as Fig. 4, but showing slices parallel to the SGX axis rather than the SGZ axis.

ratio for galaxy i given by the 2MRS or PSCz model, with the model velocity scaled according to

$$v_{i,j}/v_{i,\text{fid}} = \beta_j/\beta_{\text{fid}}. \quad (14)$$

Here, we consider a range of possible β values, β_j . $v_{i,j}$ is then the model velocity for galaxy i , assuming β value β_j . The value of β_j that minimizes χ_v^2 according to equation (13) is then our best-

fitting value of β . The 68 per cent uncertainty on β is then located at $\chi_v^2 = \chi_{v,\text{min}}^2(1 \pm 1/N)$.

The resulting best-fitting values of β are 0.16 ± 0.04 for 2MRS and 0.41 ± 0.04 for PSCz. The corresponding values of χ_v^2 are 1.14 and 1.09, respectively. These values can be compared to the values measured by other authors for the same velocity field reconstructions, as shown in Table 1. As seen there, our measured value of β for PSCz is on the low end of the literature values, though nearly

Table 1. The best-fitting values of β measured for both the 2MRS and PSCz models, using χ^2 minimization, for a variety of scenarios explained in greater detail in Section 4.2. We also list several literature values for β for each of the two models, again with 2MRS values in the left-hand column and PSCz values in the right-hand column.

2MRS type of fit	β	χ^2_v	PSCz type of fit	β	χ^2_v
Fiducial value	0.40		Fiducial value	0.50	
Standard fit	0.17 ± 0.04	1.15	Standard fit	0.41 ± 0.04	1.10
Fix $\langle \Delta d \rangle$ to 2MTF value	0.31 ± 0.04	1.10	Fix $\langle \Delta d \rangle$ to 2MTF value	0.41 ± 0.04	1.11
Fit $\langle \Delta d \rangle$	0.35 ± 0.04	1.10	Fit $\langle \Delta d \rangle$	0.40 ± 0.05	1.09
Literature values					
Branchini, Davis & Nusser (2012)	0.32 ± 0.08		Branchini et al. (2001)	0.42 ± 0.04	
Nusser, Branchini & Davis (2012)	0.32 ± 0.10		Nusser et al. (2001)	0.50 ± 0.10	
Davis et al. (2011)	0.33 ± 0.04		Zaroubi et al. (2002)	0.51 ± 0.06	
Bilicki et al. (2011)	0.38 ± 0.04		Ma, Branchini & Scott (2012)	0.53 ± 0.01	
Lavaux et al. (2010)	~ 0.52		Turnbull et al. (2012)	0.53 ± 0.08	
Pike & Hudson (2005)	0.55 ± 0.05		Radburn-Smith, Lucey & Hudson (2004)	0.55 ± 0.06	

identical to the value measured by Branchini et al. (2001). On the other hand, the value of β for 2MRS is substantially lower than any of the literature values.

As noted in Section 4.1, there is a significant monopole offset between 2MTF and PSCz. We define $\langle \Delta d \rangle$ as the mean value of the logarithmic distance ratio Δd , averaged over galaxies in either the data or the models. This gives us the zero-point of the data set in question. We find that $\langle \Delta d \rangle = -0.003$ for 2MTF, $+0.013$ for the 2MRS model, and -0.003 for the PSCz model, if measured only at the positions of the 2MTF galaxies. Thus, there is a significant monopole discrepancy between 2MTF and the 2MRS model.

In the case of both the data and the models, however, the zero-point of the velocity field is based on a set of assumptions about the boundary conditions that cannot be independently tested with the data at hand. The 2MRS model used here, for example, assumes zero gravitational potential at the survey boundary, though the authors also considered alternative sets of assumptions, such as zero net velocity at the survey boundary. The 2MTF data set likewise assumes $\langle \Delta d \rangle$ is ~ 0 in the survey volume by construction, because a monopole in the velocity field is completely degenerate with an offset in the zero-point of the TF relation. We thus consider how our measurement of β might differ with a different set of assumptions about the zero-point. If we shift all of the Δd values of 2MRS by 0.018 dex, so that its value of $\langle \Delta d \rangle$ matches that of 2MTF, then the best-fitting value of β is 0.31 ± 0.04 (also listed in Table 1), representing better agreement with the literature values. In Fig. 7, we show this adaptively smoothed version of the 2MRS and PSCz reconstructions, with the $\langle \Delta d \rangle$ set to -0.005 dex for both models, as it is in 2MTF.

Finally, we also allow the zero-point of the models to float as a free parameter, adjusting the zero-point to the value that gives us the minimum value of χ^2 . The best-fitting value of $\langle \Delta d \rangle$ is -0.010 for 2MRS and -0.014 for PSCz. This corresponds to $\beta = 0.35 \pm 0.04$ and $\beta = 0.40 \pm 0.05$, respectively. This again shows better agreement with the literature values for 2MRS, but no impact on the β value for PSCz. The large monopole offset between 2MRS and 2MTF suggests that shifting $\langle \Delta d \rangle$, either by fixing it to the 2MTF value or fitting it via χ^2 minimization, gives us a more meaningful estimate of β . Hereafter, we refer to these two approaches as the ‘fix’ and ‘fit’ cases, respectively.

Since β is related to the bias parameter b through the matter density Ω_m according to equation (10), we can use either an independently measured value of b to solve for Ω_m or an independently

measured value of Ω_m to solve for b . Beutler et al. (2012), for example, using a prior on H_0 of $73.8 \pm 2.4 \text{ km s}^{-1} \text{ Mpc}^{-1}$ as taken from Riess et al. (2011), and derives values of $\Omega_m = 0.250 \pm 0.022$ and $b = 1.48 \pm 0.27$ for the redshift survey component of 6dFGS (Jones et al. 2004, 2009). While we do not use 6dFGS results directly in this paper, the survey is responsible for the bulk of the Southern hemisphere galaxies found in 2MRS. We might then expect the surveys to have a similar value of the bias parameter. If we assume this value of $b = 1.48 \pm 0.27$ for 2MRS, and use our measured $\beta = 0.31 \pm 0.04$ from the scenario in which we fix $\langle \Delta d \rangle$ to the 2MTF value, then that gives $\Omega_m = 0.24 \pm 0.10$. Whereas if we use our measured $\beta = 0.35 \pm 0.04$ from the case in which we fit $\langle \Delta d \rangle$ using χ^2 minimization, then that gives $\Omega_m = 0.30 \pm 0.12$. Beutler also considers an H_0 prior of $67 \pm 3.2 \text{ km s}^{-1} \text{ Mpc}^{-1}$ as taken from Beutler et al. (2011), from which the authors derive $\Omega_m = 0.279 \pm 0.028$ and $b = 1.52 \pm 0.29$. Using that value of b would give us $\Omega_m = 0.25 \pm 0.10$ and $\Omega_m = 0.32 \pm 0.13$ in the ‘fix $\langle \Delta d \rangle$ to 2MTF value’ and ‘fit $\langle \Delta d \rangle$ ’ cases, respectively. Each of these estimates agrees with both the WMAP-9 yr value ($\Omega_m = 0.279 \pm 0.025$; Bennett et al. 2013) and the Planck 2015 data release value ($\Omega_m = 0.308 \pm 0.012$; The Planck Collaboration XIII 2015).

We can also invert the problem, and assume the value of Ω_m , to give us b . The Planck value of $\Omega_m = 0.308 \pm 0.012$ (The Planck Collaboration XIII 2015) gives us, for 2MRS, $b = 1.69 \pm 0.22$ and $b = 1.49 \pm 0.17$ in the ‘fixed’ and ‘fit’ cases, respectively. From this, we can derive the growth rate of structure $\Gamma \sim f\sigma_8$, where σ_8 refers to the amplitude of mass fluctuations on scales of $8 h^{-1} \text{ Mpc}$, while $f = \Omega_m^{0.55}$. This is related to $\sigma_{8,g}$, the amplitude of galaxy fluctuations on the same scale, by $\sigma_{8,g} = b\sigma_8$. We can get the value of $\sigma_{8,g}$ for the 2MRS sample from Westover (2007), who measures $\sigma_{8,g} = 0.97 \pm 0.05$. This then gives us $f\sigma_8 = \beta\sigma_{8,g} = 0.30 \pm 0.04$ and 0.34 ± 0.04 for the ‘fixed’ and ‘fit’ cases, respectively. These values are comparable to $f\sigma_8 = 0.31 \pm 0.05$, as found by Davis et al. (2011), but somewhat lower than the $f\sigma_8 = 0.42 \pm 0.07$ found by Beutler et al. (2012), the $f\sigma_8 = 0.42 \pm 0.06$ found by Beutler et al. (2012), and the $f\sigma_8 = 0.418 \pm 0.065$ found by Johnson et al. (2014). The value also falls below both the WMAP-9 yr value of 0.41 ± 0.02 and the Planck 2015 data release value of 0.44 ± 0.01 .

4.3 Residual bulk flow

Hong et al. (2014) measured the bulk flow of the 2MTF sample using a χ^2 minimization method. The authors used the following

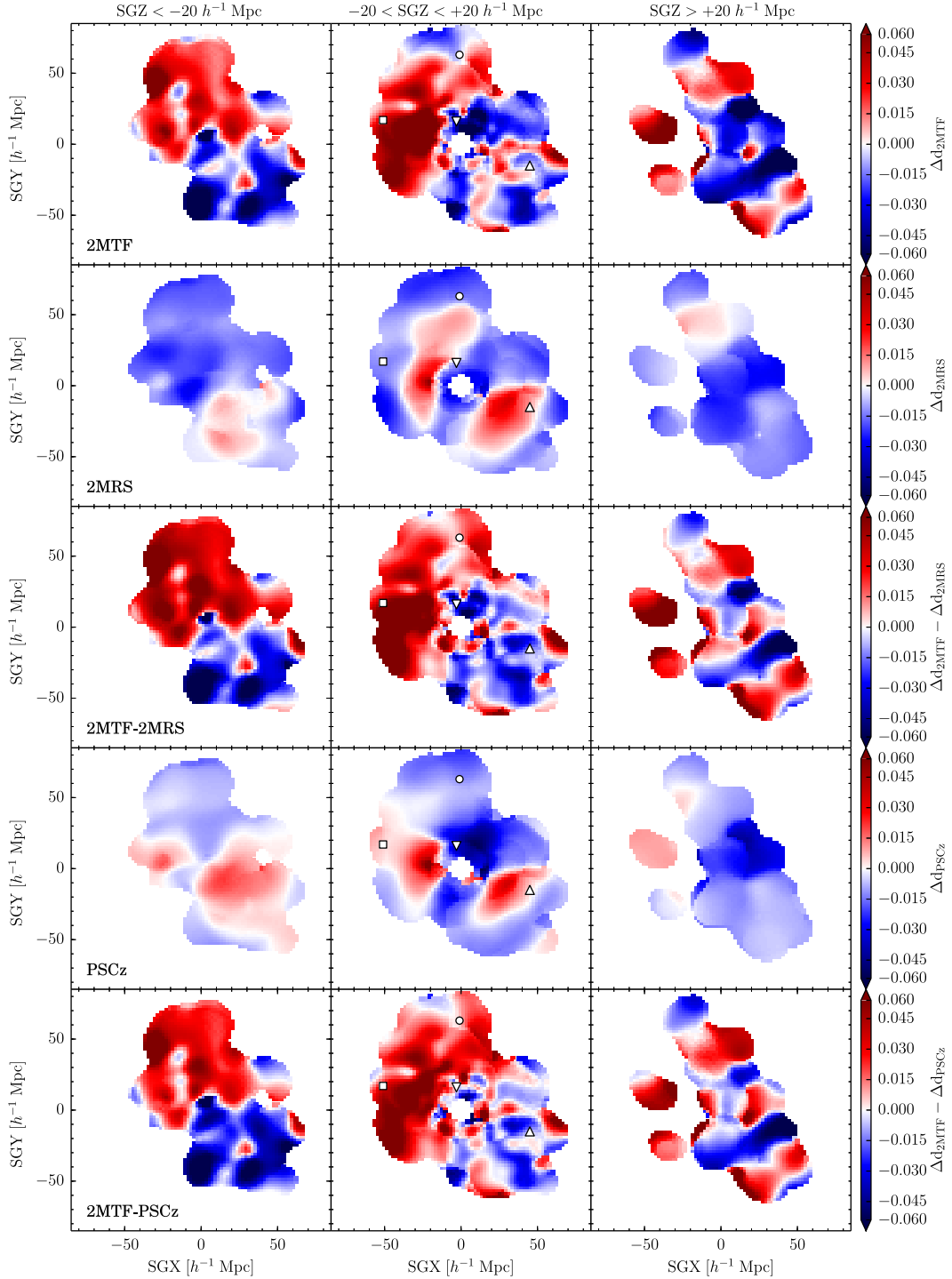


Figure 7. Same as Fig. 4, but with the $\Delta d_{2\text{MRS}}$ and Δd_{PSCz} values adjusted so that their mean value is -0.005 dex, which is the mean value of $\Delta d_{2\text{MTF}}$.

relation for χ^2 :

$$\chi^2 = \sum_{i=1}^N \frac{(\Delta d_i - \Delta d_{\text{model},i})^2 w_i^r w_i^d}{\sigma_i^2 \sum_{i=1}^N w_i^r w_i^d}, \quad (15)$$

where w_i^r is the weight assigned to galaxy i based on the radial distribution of the sample, while w_i^d is the weight that accounts for the variation in completeness as a function of declination. Note that in this case, $\Delta d_{\text{model},i}$ refers to a model in which the velocity

field is characterized by a single dipole velocity vector \mathbf{V}_{bulk} , which we refer to as the bulk flow. This is in contrast to $\Delta d_{\text{model},i}$ from equation (13) in this paper, for which the model in question is the 2MRS or PSCz reconstruction.

The radial weight w_i^r is set so that the redshift distribution of the sample is adjusted to match that of a Gaussian density profile, following Watkins, Feldman & Hudson (2009):

$$\rho(r) \propto \exp(-r^2/2R_I^2) \quad (16)$$

Table 2. Best-fitting values of the residual bulk flow, derived for Gaussian window functions of depth 20, 30, and 40 h^{-1} Mpc, using χ^2 minimization. In the top three rows, we show the total bulk flow as derived by Hong et al. (2014). The subsequent rows show residual bulk flows for the 2MRS and PSCz models under various assumptions about the best-fitting value of β , and the zero-point ($\langle \Delta d \rangle$) of the velocity field for the model in question, as described in Section 4.2. We provide the amplitude of the residual bulk flow $|V_{\text{resid}}|$, as well as its sky position, expressed in both spherical (*sgl*, *sgb*) and Cartesian supergalactic ($V_{\text{resid,sgx}}$, $V_{\text{resid,sgy}}$, $V_{\text{resid,sgz}}$) coordinates. We also translate the supergalactic coordinates into Galactic coordinates *l* and *b*. The bulk flow measurements from Hong et al. (2014) were made in Galactic coordinates, and so the measurement errors are only expressed in those coordinates. On the other hand, the residual bulk flow measurements in this paper were made in supergalactic coordinates, and so errors are given only for that case.

Type of fit	Depth (h^{-1} Mpc)	β	$\langle \Delta d \rangle$ (dex)	$ V_{\text{resid}} $ (km s^{-1})	sgl (deg)	sgb (deg)	$V_{\text{resid,sgx}}$ (km s^{-1})	$V_{\text{resid,sgy}}$ (km s^{-1})	$V_{\text{resid,sgz}}$ (km s^{-1})	<i>l</i> (deg)	<i>b</i> (deg)
Bulk flow	20	0		311 ± 34	164	−27				288 ± 6	11 ± 3
Bulk flow	30	0		281 ± 25	158	−18				296 ± 16	19 ± 6
Bulk flow	40	0		292 ± 28	171	−20				296 ± 10	6 ± 9
2MRS fiducial value	20	0.40	0.013	228 ± 14	169 ± 2	-5 ± 4	-223 ± 15	43 ± 10	-20 ± 15	311	10
2MRS fiducial value	30	0.40	0.013	271 ± 14	167 ± 2	-12 ± 3	-259 ± 13	60 ± 12	-55 ± 13	304	11
2MRS fiducial value	40	0.40	0.013	282 ± 15	165 ± 3	-18 ± 3	-258 ± 15	71 ± 13	-88 ± 13	297	12
2MRS standard fit	20	0.16	0.013	304 ± 14	167 ± 2	-10 ± 3	-292 ± 15	70 ± 10	-50 ± 16	306	12
2MRS standard fit	30	0.16	0.013	330 ± 14	166 ± 2	-15 ± 2	-309 ± 15	77 ± 12	-84 ± 12	301	10
2MRS standard fit	40	0.16	0.013	330 ± 14	165 ± 2	-20 ± 2	-299 ± 16	80 ± 14	-113 ± 11	295	10
2MRS fix $\langle \Delta d \rangle$ to 2MTF	20	0.28	−0.005	110 ± 17	182 ± 6	22 ± 10	-102 ± 17	-4 ± 10	40 ± 16	339	1
2MRS fix $\langle \Delta d \rangle$ to 2MTF	30	0.28	−0.005	165 ± 18	171 ± 4	2 ± 5	-163 ± 19	27 ± 13	4 ± 15	318	9
2MRS fix $\langle \Delta d \rangle$ to 2MTF	40	0.28	−0.005	189 ± 16	165 ± 4	-11 ± 4	-179 ± 15	49 ± 11	-36 ± 14	304	13
2MRS fit $\langle \Delta d \rangle$	20	0.35	−0.010	92 ± 20	195 ± 10	41 ± 10	-67 ± 17	-18 ± 10	60 ± 17	0	−7
2MRS fit $\langle \Delta d \rangle$	30	0.35	−0.010	136 ± 18	173 ± 5	10 ± 7	-133 ± 19	16 ± 13	23 ± 16	327	8
2MRS fit $\langle \Delta d \rangle$	40	0.35	−0.010	159 ± 16	165 ± 4	-7 ± 5	-153 ± 19	41 ± 12	-19 ± 14	309	14
PSCz fiducial value	20	0.50	−0.003	189 ± 20	159 ± 6	32 ± 7	-150 ± 19	57 ± 17	100 ± 19	349	21
PSCz fiducial value	30	0.50	−0.003	216 ± 21	156 ± 4	10 ± 4	-194 ± 20	87 ± 15	38 ± 13	325	25
PSCz fiducial value	40	0.50	−0.003	221 ± 17	154 ± 4	-2 ± 3	-198 ± 15	99 ± 17	-6 ± 12	312	26
PSCz standard fit	20	0.41	−0.003	212 ± 18	161 ± 5	19 ± 5	-189 ± 19	65 ± 15	69 ± 17	335	20
PSCz standard fit	30	0.41	−0.003	244 ± 20	158 ± 3	3 ± 3	-226 ± 18	90 ± 12	12 ± 14	329	23
PSCz standard fit	40	0.41	−0.003	248 ± 16	156 ± 3	-7 ± 3	-225 ± 15	99 ± 14	-29 ± 12	307	23
PSCz fix $\langle \Delta d \rangle$ to 2MTF	20	0.41	−0.005	211 ± 27	54 ± 83	86 ± 6	8 ± 22	11 ± 17	210 ± 26	50	10
PSCz fix $\langle \Delta d \rangle$ to 2MTF	30	0.41	−0.005	147 ± 27	138 ± 12	62 ± 8	-52 ± 23	46 ± 13	129 ± 21	25	24
PSCz fix $\langle \Delta d \rangle$ to 2MTF	40	0.41	−0.005	119 ± 26	140 ± 10	40 ± 11	-70 ± 25	59 ± 12	76 ± 17	2	34
PSCz fit $\langle \Delta d \rangle$	20	0.40	−0.014	218 ± 28	44 ± 89	86 ± 7	10 ± 23	10 ± 18	217 ± 27	50	9
PSCz fit $\langle \Delta d \rangle$	30	0.40	−0.014	150 ± 28	138 ± 13	63 ± 8	-50 ± 23	46 ± 14	134 ± 21	26	24
PSCz fit $\langle \Delta d \rangle$	40	0.40	−0.014	121 ± 27	139 ± 10	41 ± 12	-69 ± 25	59 ± 12	79 ± 18	4	34

which translates to the number distribution

$$n(r) \propto r^2 \exp(-r^2/2R_l^2), \quad (17)$$

where R_l is the characteristic depth of the bulk flow measurement.

The declination weight gives all galaxies north of $\delta = -40^\circ$ a weight of 1.00, and all galaxies south of $\delta = -40^\circ$ a weight of 2.08. This is done to account for the fact that the 2MTF survey completeness is ~ 2.08 times greater north of $\delta = -40^\circ$ than it is south of that declination, owing to the differences in sensitivity of the Northern and Southern hemisphere telescopes used in the survey.

Errors on the bulk flow were estimated using a jackknife approach, as outlined in section 4.1 of Hong et al. (2014). 50 jackknife subsamples are created, with each randomly removing 2 per cent of the 2MTF sample. The χ^2 minimization is performed separately on each subsample, and the resulting scatter is then converted into a statistical uncertainty on the bulk flow, according to Hong et al. (2014) equation (9).

The resulting bulk flow measurements at depths of 20, 30, and 40 h^{-1} Mpc are given in table 1 of Hong et al. (2014), and we reproduce those numbers in Table 2 of this paper. As noted by Hong et al. (2014), these values for the bulk flow are consistent at the 1σ level with the expectations given by the Λ CDM model (using the Λ CDM parameters $\Omega_m = 0.27$, $\Omega_\Lambda = 0.73$, $n_s = 0.96$, as taken from the WMAP-7 yr results of Larson et al. 2011, and the

matter power spectrum generated by the CAMB package, as given by Lewis, Challinor & Lasenby 2000).

Separate from the question of whether the bulk flow agrees with predictions from Λ CDM, there is also the question of whether the observed amplitude and direction of the bulk flow is predicted by the particular galaxy density distribution we observe in the local Universe. We now investigate a scenario in which the $\Delta d_{\text{model},i}$ from equation (15) is calculated using the model velocity of the 2MRS or PSCz models plus a residual bulk flow V_{resid} . Fixing β to the fiducial values of 0.40 and 0.50, respectively, and performing the χ^2 minimization as in Hong et al. (2014), we get residual bulk flow values of amplitude $\sim 200 \text{ km s}^{-1}$ at depths of 20, 30, and 40 h^{-1} Mpc for PSCz, with somewhat larger values for 2MRS.

In addition to the fiducial values of β , we also calculate the residual bulk flow measured for the fitted values of β using the methods described in Section 4.2. Each of these values are listed in Table 2. (Note that, for ease of comparison with previous papers on galaxy peculiar velocities and large-scale structure, we have done the fitting in supergalactic coordinates, whereas Hong et al. (2014) fit the bulk flow in Galactic coordinates.) For both models, the amplitude of the residual bulk flow in the ‘standard fit’ case is seen to be larger than in the fiducial case. In fact, the 2MRS residual bulk flow for the standard fit of β is even larger in amplitude than the total bulk flow. As we noted in Section 4.2, though, we consider the ‘fix’ and ‘fit’ scenarios to offer more realistic estimates of the

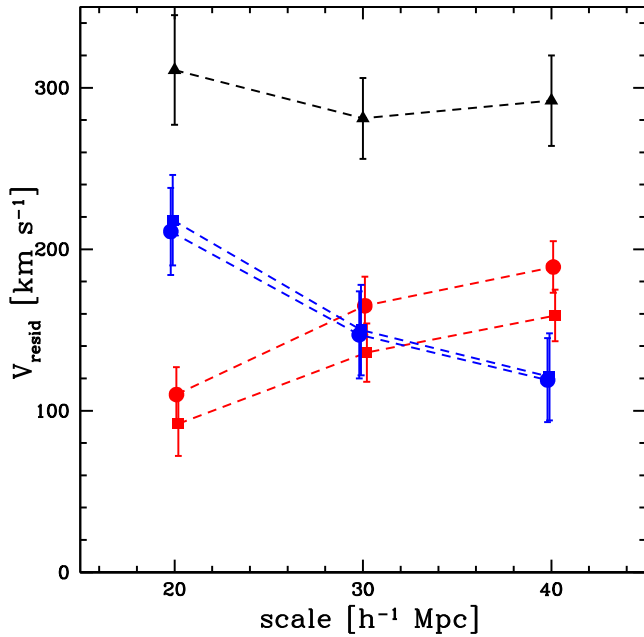


Figure 8. Black triangles show the magnitude of the 2MTF bulk flow with Gaussian window functions at depths of 20, 30, and 40 h^{-1} Mpc, as derived by Hong et al. (2014). Red (blue) points represent residual bulk flow amplitudes at the same depths for the 2MRS (PSCz) model, with circles showing the values for the ‘fix (Δd) to 2MTF value’ scenario, and squares for the ‘fit (Δd)’ scenario. While each of the points lies exactly at either 20, 30, or 40 h^{-1} Mpc, we give a slight horizontal displacement to the points in order to make it more clear which set of error bars corresponds to which data point.

underlying value of β . For both of these fits, for both the 2MRS and PSCz models, the residual bulk flow is $\sim 150 \text{ km s}^{-1}$, which corresponds to roughly half of the total bulk flow. This is also illustrated in Fig. 8.

The residual bulk flow directions in both the ‘fiducial’ and ‘standard fit’ cases lie close to that of the *total* bulk flow direction. In Fig. 9, however, we show the residual bulk flow directions in both the ‘fix’ and ‘fit’ cases. In both cases, the 2MRS residual extends from very low Galactic latitudes at the 20 h^{-1} Mpc scale to somewhat higher latitudes (and a direction very close to both the total bulk flow and the Hydra-Centaurus Supercluster) at the 40 h^{-1} Mpc scale. The PSCz residual bulk flow in the ‘fix’ and ‘fit’ cases follow a similar pattern on the sky, but offset by $\sim 60^\circ$. The PSCz residual at 40 h^{-1} Mpc is $\sim 45^\circ$ away from the Shapley Supercluster, and otherwise does not appear to be in the vicinity of any major features of large-scale structure.

Figs 4–7 also show the residual bulk flow cosmographically, in the $\Delta d_{2\text{MTF}} - \Delta d_{2\text{MRS}}$ and $\Delta d_{2\text{MTF}} - \Delta d_{\text{PSCz}}$ rows. These can be compared to figs 9 and 10 from Springob et al. (2014), which show the cosmography of the 6dFGS peculiar velocity field, when the 2MRS and PSCz models are subtracted away from the observed Δd values. 6dFGS is a Southern hemisphere only survey, and is composed of galaxies whose mean redshift is more than twice as great as that of 2MTF. None the less, Springob et al. (2014) finds that the PSCz model offers a somewhat better fit to the 6dFGS velocity field than does the 2MRS model. Cosmographically, Springob et al. (2014) figs 9 and 10 show that the models underestimate both the outward flow towards Hydra-Centaurus and the Shapley Supercluster *and* the inward flow coming from a direction on nearly the opposite end of the sky, roughly coincident with the Cetus Super-

cluster. In this paper, however, comparing the 2MTF velocity field to that of PSCz, we find that while the model underestimates the flow towards the Shapley Supercluster, there is not such a large underestimate of the inflow from the anti-Shapley direction. (See the bottom central panel of Fig. 4.) Thus, at least judging from 2MTF, it seems possible that the deficiency of the models is that they underestimate the impact of known structures such as Shapley and Hydra-Centaurus, rather than that there are large structures outside the survey volume with a large influence on the local velocity field.

We can compare the residual bulk flow direction to results from other authors. Hudson et al. (2004) compares peculiar velocities from the ‘Streaming Motions of Abell Clusters’ (Hudson et al. 2001) to the PSCz model, and finds a residual bulk flow of $372 \pm 127 \text{ km s}^{-1}$ towards $(l, b) = (273^\circ, 6^\circ)$, which is offset from our PSCz residual bulk flow by $\sim 90^\circ$. Magoulas (2012) compares 6dFGS peculiar velocities to the predictions of the 2MRS model, and finds a residual bulk flow of 273 ± 45 towards $(l, b) = (326^\circ, 37^\circ)$. This is $\sim 30^\circ$ offset from our 2MRS residual bulk flow, but similarly close to the Shapley Supercluster. Carrick et al. (2015) compares the SFI++ (Springob et al. 2007) TF and First Amendment (Turnbull et al. 2012) Type Ia Supernovae peculiar velocities to the predictions of the 2M++ reconstruction (Lavaux & Hudson 2011), finding a residual bulk flow of $159 \pm 23 \text{ km s}^{-1}$ towards $(l, b) = (304^\circ, 6^\circ)$ at 50 h^{-1} Mpc, very close to our own measured 2MRS residual bulk flow at 40 h^{-1} Mpc in both amplitude and direction. There is heavy overlap between the 2MRS and 2M++ samples, so this agreement should not be surprising.

How does the amplitude of the residual bulk flow compare to our expectations, given a standard Λ CDM framework? Hudson et al. (2004) examined the question of how consistent the rest frame of the PSCz gravity field should be with the CMB frame. They estimated that while the contribution to the local bulk motion from sources beyond $\sim 200 h^{-1}$ Mpc should only be $\sim 50 \text{ km s}^{-1}$, the contribution from systematic uncertainties ($\sim 90 \text{ km s}^{-1}$) and shot noise ($\sim 70 \text{ km s}^{-1}$) suggest a total uncertainty in PSCz’s reconstruction of the bulk flow of $\sim 150 \text{ km s}^{-1}$. Nusser, Davis & Branchini (2014) reached a similar conclusion with regard to 2MRS. Both Davis et al. (2011) and Carrick et al. (2015) find agreement between the measured SFI++ bulk flow and the predictions of models roughly within this expected $\sim 150 \text{ km s}^{-1}$ range. In this paper, we similarly find a similar residual bulk flow of $\sim 150 \text{ km s}^{-1}$ with respect to both the 2MRS and PSCz models, but only if one adjusts the comparison between data and model so as to remove the monopole deviation. The amplitude of the residual bulk flow appears to be heavily dependent on assumptions about the boundary conditions of the model velocity field.

5 CONCLUSIONS

We have used adaptive kernel smoothing to present a cosmographic view of the local peculiar velocity field, as measured by the TF peculiar velocities from the 2MTF survey. By extending all the way down to Galactic latitudes of $|b| = 5^\circ$, 2MTF presents a more complete view of the velocity field at $cz < 10000 \text{ km s}^{-1}$ than has been seen before.

We compare the 2MTF velocity field to the reconstructed velocity field models from the 2MRS and PSCz redshift surveys. We find best-fitting values of β for these two models of 0.17 ± 0.04 and 0.41 ± 0.04 , respectively, with χ^2_{ν} values of 1.15 and 1.10, respectively. There is a significant monopole offset between the 2MTF data set and the 2MRS model. As we have noted here, the zero-point, and resulting monopole term in the velocity field, for both the data and

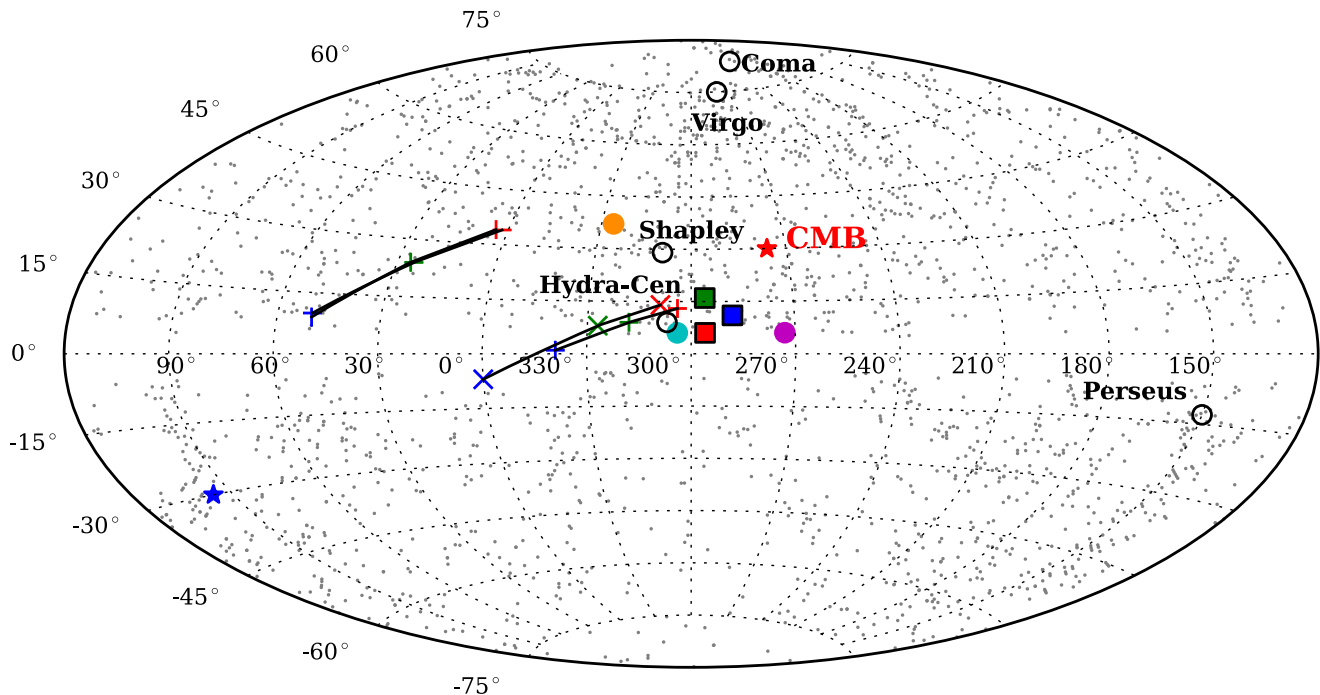


Figure 9. The direction of the residual bulk flow of the 2MTF velocity field, as projected on the celestial sphere in Galactic coordinates, with respect to the 2MRS and PSCz models. The residual bulk flow fit for 20, 30, and 40 h^{-1} Mpc window functions are shown in blue, green, and red, respectively, with black lines connecting them. 2MRS residual bulk flow shown as +’s for the ‘fix’ case and x’-s for the ‘fit’ case. PSCz residual bulk flows are shown as -’s and |’s in the fix and fit cases, respectively. As also shown in Table 2, the fix and fit directions for PSCz are nearly identical. We also show the total bulk flow as squares, coded with the same blue/green/red scheme for the 20, 30, and 40 h^{-1} Mpc window functions. The Carrick et al. (2015), Hudson et al. (2004), and Magoulas (2012) residual bulk flows from the 2M++, PSCz, and 2MRS models, respectively, are shown as cyan, magenta, and orange circles. The red star shows the location of the CMB dipole direction, with the blue star representing the antipole direction. Grey dots show the locations of 2MTF galaxies. The approximate positions of various features of large-scale structure are shown as open circles with labels.

models has been set on the basis of assumptions about the boundary conditions. By changing those assumptions we may investigate the impact of the zero-point on the measurement of β , as well as other parameters. If we remove any monopole offset between data and models for both 2MRS and PSCz, then the best-fitting values of β become 0.31 ± 0.04 and 0.41 ± 0.04 , with χ^2_v . Fitting the zero-point to minimize χ^2_v yields similar results.

These latter β values are in line with previous estimates of β for these same models from the literature. They can be used in conjunction with other measurements to estimate parameters such as the matter density Ω_m and the growth rate of structure $f\sigma_8$. We estimate values of Ω_m in line with previous measurements. We also find a growth rate of structure of ~ 0.3 , in line with the estimate by Davis et al. (2011), but somewhat lower than most other estimates.

Hong et al. (2014) used a χ^2 minimization technique with Gaussian window function at depths of 20, 30, and 40 h^{-1} Mpc to measure the bulk flow of the 2MTF velocity field. In this paper, we have now used that same technique to measure the *residual* bulk flow: the component of the bulk flow not accounted for by the velocity field models. If we assume either the fiducial value of β or the standard χ^2 minimization fit, then the residual bulk flow is $\sim 200\text{--}300$ km s^{-1} , of roughly the same amplitude and direction as the total bulk flow.

However, when we adjust the zero-point of the models, either by fixing them to the 2MTF value or fitting them via χ^2 minimization, then the residual bulk flow drops to ~ 150 km s^{-1} . This is in agreement with theoretical expectations, as discussed by Hudson et al. (2004) and Nusser, et al. (2014). The residual bulk flow direction for 2MRS is at low Galactic latitude, and within $\sim 15^\circ$ of both Hydra-

Centaurus and the total bulk flow direction. On the other hand, the PSCz residual direction is at higher Galactic latitude (34° at the 40 h^{-1} Mpc scale) and not associated with any known features of large-scale structure. Carrick et al. (2015) finds a residual bulk flow comparing their TF and SNe peculiar velocity sample to the 2M++ sample which is strikingly similar in both amplitude and direction to our measured residual bulk flow for the 2MRS reconstruction.

This suggests that while 2MTF extends towards lower Galactic latitudes and creates a more complete all-sky sample, it does not reveal any significant influence on the velocity field from hidden structures not already apparent from peculiar velocity surveys with a larger zone of avoidance, like SFI++. Both the total bulk flow and the residual bulk flow from the 2MRS model are at low Galactic latitude, but that was already observed in earlier surveys. The inclusion of the lower Galactic latitude galaxies does not appear to create a substantial shift in the residual bulk flow direction.

ACKNOWLEDGEMENTS

The authors gratefully acknowledge Martha Haynes, Riccardo Giovanelli, and the ALFALFA team for supplying the latest ALFALFA survey data. We also thank Enzo Branchini for providing a copy of the PSCz reconstruction. We acknowledge the contributions of John Huchra (1948–2010) in the development of the 2MRS and 2MTF samples used in this paper.

This research was conducted by the Australian Research Council Centre of Excellence for All-sky Astrophysics (CAASTRO), through project number CE110001020.

REFERENCES

- Aaronson M., Huchra J., Mould J., Schechter P. L., Tully R. B., 1982, *ApJ*, 258, 64
- Bennett C. L. et al., 2013, *ApJS*, 208, 20
- Beutler F. et al., 2011, *MNRAS*, 416, 3017
- Beutler F. et al., 2012, *MNRAS*, 423, 3430
- Bilicki M., Chodorowski M., Jarrett T., Mamon G. A., 2011, *ApJ*, 741, 31
- Branchini E. et al., 1999, *MNRAS*, 308, 1
- Branchini E. et al., 2001, *MNRAS*, 326, 1191
- Branchini E., Davis M., Nusser A., 2012, *MNRAS*, 424, 472
- Carrick J., Turnbull S. J., Lavaux G., Hudson M. J., 2015, *MNRAS*, 450, 317
- Crook A. C., Huchra J. P., Martimbeau N., Masters K. L., Jarrett T., Macri L. M., 2007, *ApJ*, 655, 790
- Davis M., Nusser A., Willick J. A., 1996, *ApJ*, 473, 22
- Davis M., Nusser A., Masters K. L., Springob C. M., Huchra J. P., Lemson G., 2011, *MNRAS*, 413, 2906
- Djorgovski S., Davis M., 1987, *ApJ*, 313, 59
- Dressler A., Lynden-Bell D., Burstein D., Davies R. L., Faber S. M., Terlevich R., Wegner G., 1987, *ApJ*, 313, 42
- Ebeling H., White D. A., Rangarajan F. V. N., 2006, *MNRAS*, 368, 65
- Efstathiou G., Ellis R. S., Peterson B. A., 1988, *MNRAS*, 232, 431
- Erdođdu P. et al., 2006, *MNRAS*, 373, 45
- Fisher K. B., Lahav O., Hoffman Y., Lynden-Bell D., Zaroubi S., 1995, *MNRAS*, 272, 885
- Giovanelli R., Haynes M. P., Salzer J. J., Wegner G., da Costa L. N., Freudling W., 1994, *AJ*, 107, 2036
- Giovanelli R., Haynes M. P., Salzer J. J., Wegner G., da Costa L. N., Freudling W., 1995, *AJ*, 110, 1059
- Giovanelli R., Haynes M. P., Herter T., Vogt N. P., Wegner G., Salzer J. J., da Costa L. N., Freudling W., 1997a, *AJ*, 113, 22
- Giovanelli R., Haynes M. P., Herter T., Vogt N. P., da Costa L. N., Freudling W., Salzer J. J., Wegner G., 1997b, *AJ*, 113, 53
- Giovanelli R. et al., 2005, *AJ*, 130, 2598
- Harrison E. R., 1974, *ApJ*, 191, L51
- Haynes M. P., Giovanelli R., Salzer J. J., Wegner G., Freudling W., da Costa L. N., Herter T., Vogt N. P., 1999a, *AJ*, 117, 1668
- Haynes M. P., Giovanelli R., Chamaroux P., da Costa L. N., Freudling W., Salzer J. J., Wegner G., 1999b, *AJ*, 117, 2039
- Haynes M. P. et al., 2011, *AJ*, 142, 170
- Hoffman Y., Courtois H. M., Tully R. B., 2015, *MNRAS*, 449, 4494
- Hong T. et al., 2013, *MNRAS*, 432, 1178
- Hong T. et al., 2014, *MNRAS*, 445, 402
- Huchra J. et al., 2005, in Fairall K. P., Woudt P. A., eds, *ASP Conf. Ser. Vol. 329, Nerby Large-Scale Structures and the Zone of Avoidance*. Astron. Soc. Pac., San Francisco, p. 135
- Huchra J. P. et al., 2012, *ApJS*, 199, 26
- Hudson M. J., 1994, *MNRAS*, 266, 475
- Hudson M. J., Lucey J. R., Smith R. J., Schlegel D. J., Davies R. L., 2001, *MNRAS*, 327, 265
- Hudson M. J., Smith R. J., Lucey J. R., Branchini E., 2004, *MNRAS*, 352, 61
- Jarrett T. H., Chester T., Cutri R., Schneider S., Skrutskie M., Huchra J. P., 2000, *AJ*, 119, 2498
- Johnson A. et al., 2014, *MNRAS*, 444, 3926
- Jones D. H. et al., 2004, *MNRAS*, 355, 747
- Jones D. H., Peterson B. A., Colless M., Saunders W., 2006, *MNRAS*, 369, 25
- Jones D. H. et al., 2009, *MNRAS*, 399, 683
- Kaiser N., Efstathiou G., Saunders W., Ellis R., Frenk C., Lawrence A., Rowan-Robinson M., 1991, *MNRAS*, 252, 1
- Kochanek C. S. et al., 2001, *ApJ*, 560, 566
- Larson D. et al., 2011, *ApJS*, 192, 16
- Lavaux G., Hudson M. J., 2011, *MNRAS*, 416, 2840
- Lavaux G., Tully R. B., Moyahee R., Colombi S., 2010, *ApJ*, 709, 483
- Lewis A., Challinor A., Lasenby A., 2000, *ApJ*, 538, 473
- Linder E. V., 2005, *Phys. Rev. D*, 72, 043529
- Lynden-Bell D., Faber S. M., Burstein D., Davies R. L., Dressler A., Terlevich R. J., Wegner G., 1988, *ApJ*, 326, 19
- Magoulas C., 2012, PhD thesis, Univ. Melbourne
- Ma Y.-Z., Branchini E., Scott D., 2012, *MNRAS*, 425, 2880
- Malmquist K. G., 1924, *Medd. Lund Astron. Obs. Ser. II*, 32, 64
- Masters K. L., Springob C. M., Haynes M. P., Giovanelli R., 2006, *ApJ*, 653, 861
- Masters K. L., Springob C. M., Huchra J. P., 2008, *AJ* 135, 1738
- Masters K. L., Springob C. M., Huchra J. P., 2014a, *AJ* 147, 124
- Masters K. L., Crook A., Hong T., Jarrett T. H., Koribalski B., Macri L., Springob C. M., Staveley-Smith L., 2014b, *MNRAS*, 443, 1044
- Matthewson D. S., Ford V. L., Buchhorn M., 1992, *ApJS*, 81, 413
- Nusser A., Davis M., 1994, *ApJ*, 421, L1
- Nusser A., da Costa L. N., Branchini E., Bernardi M., Alonso M. V., Wegner G., Willmer C. N. A., Pellegrini P. S., 2001, *MNRAS*, 320, L21
- Nusser A., Branchini E., Davis M., 2012, *ApJ*, 744, 193
- Nusser A., Davis M., Branchini E., 2014, *ApJ*, 788, 157
- Paturel G., Theureau G., Bottinelli L., Gouguenheim L., Coudreau-Durand N., Hallet N., Petit C., 2003, *A&A*, 412, 57
- Pike R. W., Hudson M. J., 2005, *ApJ*, 635, 11
- Press W. H., Schechter P., 1974, *ApJ*, 187, 425
- Radburn-Smith D. J., Lucey J. R., Hudson M. J., 2004, *MNRAS*, 355, 1378
- Riess A. G. et al., 2011, *ApJ*, 732, 129
- Saunders W. et al., 2000, *MNRAS*, 317, 55
- Shaya E. J., Tully R. B., Pierce M. J., 1992, *ApJ*, 391, 16
- Silverman B. W., 1986, *Density Estimation for Statistics and Data Analysis*. Chapman & Hall, London
- Springob C. M., Haynes M. P., Giovanelli R., Kent B., 2005, *ApJS*, 160, 149
- Springob C. M., Masters K. L., Haynes M. P., Giovanelli R., Marinoni C., 2007, *ApJS*, 172, 599
- Springob C. M. et al., 2014, *MNRAS*, 445, 4267
- Strauss M. A., Willick J. A., 1995, *Phys. Rep.* 261, 271
- The Planck Collaboration XIII, 2015, preprint ([astro-ph/1502.01589](https://arxiv.org/abs/astro-ph/1502.01589))
- Theureau G., Bottinelli L., Coudreau-Durand N., Gouguenheim L., Hallet N., Loulergue M., Paturel G., Teerikorpi P., 1998, *A&AS*, 130, 333
- Theureau G. et al., 2005, *A&A*, 430, 373
- Theureau G., Hanski M. O., Coudreau N., Hallet N., Martin J.-M., 2007, *A&A*, 465, 71
- Tully R. B., Fisher J. R., 1977, *A&A*, 54, 661
- Tully R. B. et al., 2013, *AJ*, 146, 86
- Tully R. B., Courtois H., Hoffman Y., Pomaredo D., 2014, *Nature* 7516, 71
- Turnbull S. J., Hudson M. J., Feldman H. A., Hicken M., Kirshner R. P., Watkins R., 2012, *MNRAS*, 420, 447
- Watkins R., Feldman H. A., Hudson M. J., 2009, *MNRAS*, 392, 743
- Westover M., 2007, PhD thesis, Harvard University
- Willick J. A., Courteau S., Faber S. M., Burstein D., Dekel A., 1995, *ApJ*, 446, 12
- Willick J. A., Courteau S., Faber S. M., Burstein D., Dekel A., Kolatt T., 1996, *ApJ*, 457, 460
- Yahil A., Strauss M. A., Davis M., Huchra J. P., 1991, *ApJ*, 372, 380
- Zaroubi S., Branchini E., Hoffman Y., da Costa L. N., 2002, *MNRAS*, 336, 1234

This paper has been typeset from a $\text{\TeX}/\text{\LaTeX}$ file prepared by the author.